

Periodic H α variations in GL 581: Further evidence for an activity origin to GL 581d

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ABSTRACT

Context. Radial velocity measurements initially showed evidence that the M dwarf GL 581 might host up to six planetary companions. Two of these, GL 581g and GL 581d had orbital distances in the so-called “habitable zone” of the star. The existence of both of these planets have been called into question. Additional radial velocity measurements for GL 581g could not confirm its presence. A study of H α in GL 581 showed that changes in this activity indicator correlated with radial velocity variations attributed to GL 581d. Thus two planets that were important for studies of habitable planets may be artifacts of stellar activity.

Aims. Previous investigations analyzing the same activity data have reached different conclusions regarding the existence of GL 581d. We therefore investigated the H α variations for GL 581 to assess the nature of the radial velocity variations attributed to the possible planet GL 581d.

Methods. We performed a Fourier analysis of the published H α measurements for GL 581d. Fourier components are selectively found and removed in a so-called pre-whitening process thus isolating any variations at the orbital frequency of GL 581d.

Results. The frequency analysis yields five significant frequencies, one of which is associated with the 66.7 d orbital period of the presumed planet GL 581d. The H α variations at this period show sine-like variations that are 180° out-of-phase with the radial velocity variations of GL 581d. This is seen in the full data set that spans almost 7 years, as well as a subset of the data near the end of the time series that had good temporal sampling over 230 days. No significant temporal variations are found in the ratio of the amplitudes of the H α index and radial velocity variations. This provides additional evidence that the radial velocity signal attributed to GL 581d is in fact due to stellar activity.

Conclusions. The analysis confirms the anti-correlation of the radial velocity of GL 581d with the H α equivalent width and provides additional strong evidence that the signal of GL 581d is intrinsic to the star.

Key words. star: individual: GL 581 - techniques: radial velocities - stars: late-type - planetary systems

1. Introduction

Radial velocity (RV) measurements have demonstrated great successful at discovering extrasolar planets. New instrumentation, improved calibration methods, and innovative analysis techniques have steadily improved the RV precision to the point that we can routinely make measurements with precisions of $\sim 1 \text{ m s}^{-1}$ or better. At this level of precision the stellar intrinsic noise now represents a significant contribution to the measurement “error”, often referred to as the stellar RV “jitter”. In the best case this RV jitter can hinder the detection of planetary companions, in the worse case it can create a periodic signal that is interpreted as arising from a planetary companion.

M dwarf stars are objects that have become particularly attractive for Doppler surveys because the RV amplitude of the host star caused by an orbiting terrestrial planet in the habitable zone is $\sim \text{few m s}^{-1}$, a value easily measured by current techniques. Unfortunately, M dwarfs can be active and the orbital period of a planet in the habitable zone is days to weeks and this is comparable to the time scales

of stellar activity (rotational modulation, spot evolution, etc.). Activity-related RV jitter may produce false planets.

A case in point is the planetary system around the M dwarf star GL 581. Mayor et al. (2009) reported four planet in the system and shortly afterwards Vogt et al. (2010) claimed six planet candidates with periods up to 433 d. Two of these planets, GL 581d and GL 581g, were of particular interest because they had orbital distances that placed them within the habitable zone of the star. Subsequent RV measurements (Forveille et al. 2009) could not confirm the presence of GL 581g, although this has been the subject of debate (Vogt et al. 2012). Hatzes (2014) demonstrated that the signal attributed to GL 581g was real and significant, but most likely due to activity as it was not coherent on long time scales.

Baluev (2012) was the first to question the reality of GL 581d based on an analysis of the red noise in the RV data. This doubt was considerably strengthened through a study by Robertson et al. (2014). They argued that the 66-d orbital period of GL 581d was actually a harmonic of the 130 d rotation period of the star. This was based on an apparent correlation between the RV variations due to GL 581d and the equivalent width of H α , their so-called $I_{H\alpha}$ index. Correcting the RVs for this correlation eliminated the sig-

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nals attributed to GL 581d and GL 581g, while boosting the significance of the other three planets.

The nature of the 66-d RV period of GL 581 (hereafter referred to as the “orbital period” regardless of whether the planet exists or not) still remains the subject of debate. Anglada-Escudé & Tuomi (2015) questioned the conclusions of Robertson et al. (2014), arguing that their result came from an improper use of periodograms on the residual data. Although Anglada-Escudé & Tuomi agreed that there was a substantial correlation between the RV variations and $I_{H\alpha}$ index, they argued that there was no clear evidence of time variability of this index.

It is important to establish the true nature of the RV variations of the purported planet GL 581d as this has important consequences for RV searches for planets in the habitable zone of M dwarf stars. This is especially true since different analyses of the same $I_{H\alpha}$ index measurements arrived at different conclusions. Therefore, we investigated the H α variations using a different approach to the one taken by Robertson et al. (2014), namely a Fourier analysis of the frequency components in the H α time series. The aim of this work is to confirm or refute the anti-correlation between RV and H α equivalent width found by Robertson et al. (2014).

2. The Temporal Variations of H α

2.1. Frequency Analysis

For this study we used the $I_{H\alpha}$ index measurements of GL 581 from Robertson et al. (2014) which are shown in Figure 1. The data were analyzed using a Fourier approach. We first found the dominant component in the Fourier transform, and this was fit with a sine function of the appropriate period, amplitude, and phase. The contribution of this frequency was then removed from the time series and we proceeded to the next dominant sine term in the residuals. This process is often referred to as “pre-whitening”. The advantage of this technique is that since you are fitting and removing a sine function sampled in the same way as the data, alias frequencies are also removed.

Normally one associates pre-whitening with finding strictly periodic signals in your data, often in cases associated with stellar oscillations. However, pre-whitening can be more versatile in that it finds the dominant Fourier components that describe the overall variations in your time series, even if they do not appear to be periodic. The mathematical foundation for this is that trigonometric functions form basis set and a linear combination of these provide an alternative mathematical representation of the series. Pre-whitening finds those Fourier components in the time series that are clearly above the noise level. The sum of these may provide an adequate representation of the time variations in the RV due to activity.

For our pre-whitening analysis we used the program *Period04* (Lenz & Breger 2005). Table 1 lists the dominant frequencies (labeled $f_1 - f_5$), equivalent periods, and amplitudes found by the Fourier analysis. The pre-whitening procedure was stopped when the dominant peak in the residuals had a peak less than four times the surrounding noise level. Peaks of this amplitude generally have a false alarm probability of $\sim 1\%$ Kuschnig et al. (1997). Once the dominant frequency components were found a simultaneous fit was made to the data optimizing the amplitude, period,

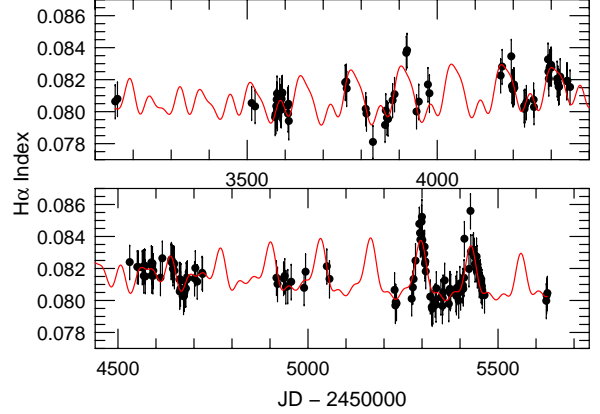


Fig. 1. Time series of the H α equivalent width measurements for GL 581. The line represents the five frequency fit using the entries in Table 1 ($f_1 - f_5$).

and phase of the individual sine terms. The multi-sine fit using these frequencies is shown as the curve in Figure 1. The first two frequencies, f_1 and f_2 are associated with the ~ 130 d rotation period of GL 581. The third frequency is the orbital frequency of GL 581d.

The orbital frequency can also be easily detected in a subset of the H α data taken JD = 2455200 and 2455645. Because of the shorter data set, fewer Fourier components were found, namely one at $\nu = 0.0079$ d $^{-1}$ ($P = 126.9$ d) and another at the orbital frequency of GL 581d, $\nu = 0.0150$ d $^{-1}$ ($P = 66.4$ d). Figure 2 shows the pre-whitening procedure on the subset $I_{H\alpha}$ data. The removal of the dominant peak at $\nu = 0.0077$ d $^{-1}$ ($P = 129$ d) results a strong peak at $\nu = 0.0152$ d $^{-1}$, or $P = 65.8$ d, near the orbital period of GL 581d.

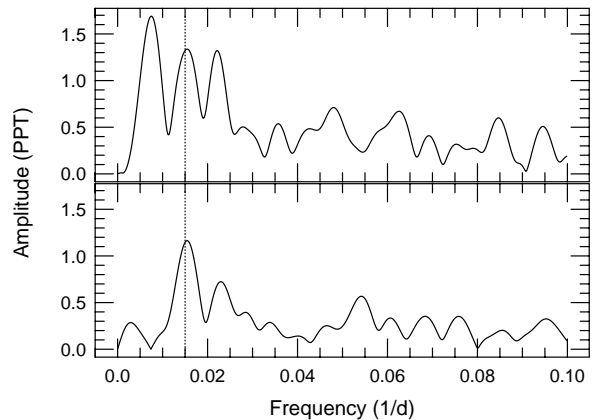


Fig. 2. Pre-whitening procedure on the H α data taken between JD = 2455200 and 2455645. The original amplitude spectrum before (top) and after (bottom) removing the dominant frequency at 0.0078 d $^{-1}$. The vertical dashed line marks the orbital frequency of GL 581d ($\nu = 0.015$ d $^{-1}$, $P = 66.6$ d).

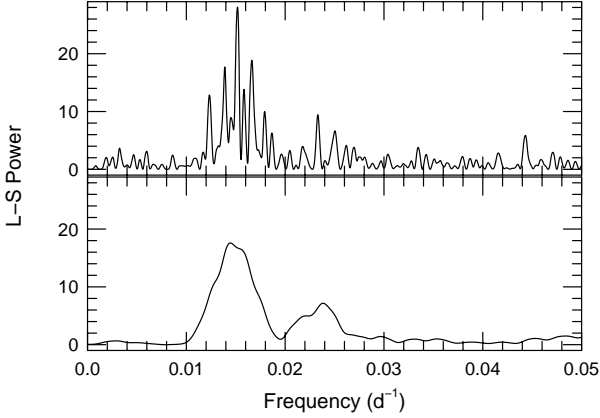


Fig. 3. (top) The L-S periodogram of all $I_{H\alpha}$ measurements after removing the contribution of all frequencies ($f_1, f_3 - f_5$ in Table 1) except that associated with GL 581d (f_2). (bottom) The L-S periodogram of the $I_{H\alpha}$ measurements taken between JD = 2455200 and 2455645 after removing the contribution of f_6 in Table 1.

2.2. Statistical Significance of the Variations

The statistical significance of a periodic signal can be assessed using the Lomb-Scargle (L-S) periodogram (Lomb 1976; Scargle 1983). The top panel in Figure 3 shows the L-S periodogram of the residual $I_{H\alpha}$ measurements after removing all frequencies except that associated with GL 581d (i.e. f_1, f_2, f_4 , and f_5 were removed). (Henceforth, we shall refer to “residual” $I_{H\alpha}$ measurements as those values where the contribution of all frequencies except f_3 have been removed from the data.)

One can estimate the false alarm probability (FAP) from the L-S power, z , and the expression $\text{FAP} = 1 - (1 - e^{-z})^N \approx Ne^{-z}$, where N is the number of independent frequencies (Scargle 1983). For the entire data set this results in $\text{FAP} \approx 10^{-13}$. For the subset data this is a $\text{FAP} \approx 10^{-6}$. The latter value was confirmed using a bootstrap procedure (Murdoch et al. 1993; Kürster et al. 1997) and 2×10^5 random shuffles of the data. In no instance was the L-S power of the random data greater than the actual value ($\text{FAP} < 5 \times 10^6$).

The L-S periodogram of the residual $I_{H\alpha}$ data from the subset data (JD = 2455200 – 2455645) is shown in the lower panel of Fig. 3. Note the dramatic increase in the L-S power when using the full data set. This indicates a long-lived and reasonably coherent signal. However, one should be wary of determining the FAP of a time series that has been modified. In this case the contributions of four frequencies were removed from the data before estimating the FAP. If one removes the dominant peaks in the periodogram the remaining ones will always look more significant. Therefore we took additional approaches to estimate the FAP.

The L-S periodogram of the original time series before pre-whitening shows L-S power of $z \approx 11.9$ at the orbital frequency of GL 581d ($\nu = 0.015 \text{ d}^{-1}$). The above expression gives us the FAP for random noise producing more power than the actual data over a broad frequency range. However, we are interested in assessing the FAP at a *known* frequency in the data, i.e. the orbital frequency of GL 581d.

In this case the FAP is given by $\text{FAP} = e^{-z}$ (Scargle 1983), or the previous expression with only one independent frequency ($N = 1$). This results in $\text{FAP} \approx 10^{-5}$.

The FAP was assessed using a revised version of the bootstrap method. The five frequency components in Table 1 were removed from the full $I_{H\alpha}$ index data. The resulting residuals represent our “noise” model for the data. The values from this noise model were then randomly shuffled keeping the times fixed. The sine functions from the frequency analysis were then added back into the noise data, except for the contribution from GL 581d, f_3 (i.e. only f_1, f_2, f_4 , and f_5 were added back in). A L-S periodogram was calculated and the maximum power in the frequency range $\nu = 0.01 - 0.02 \text{ d}^{-1}$ found. We chose this frequency range as we are interested in the probability noise would produce significant power at $\nu = 0.015 \text{ d}^{-1}$. With 2×10^5 random shuffles of the data there was no instance when the random L-S periodogram produced power higher than the original data.

We also performed the same procedure on the subset of the $I_{H\alpha}$ index data. In this case there are fewer Fourier components in the time series, namely the rotational frequency at $\nu = 0.0079 \text{ d}^{-1}$ and the orbital frequency of GL 581d, $\nu = 0.0150 \text{ d}^{-1}$. Both of these components were subtracted from the data to produce the noise model for the bootstrap procedure. As before, the orbital frequency of GL 581d ($\nu = 0.0150 \text{ d}^{-1}$) was not added back into the noise model prior to computing the L-S periodogram. Again, after 2×10^5 shuffles there was no instance of the random periodogram having higher power than the actual data. A bootstrap analysis of the full and subset data indicates that the FAP for the $I_{H\alpha}$ variations at the orbital frequency of GL 581d is $< 2 \times 10^{-6}$.

The FAP can also be estimated directly from the Fourier amplitude spectrum and the height of a peak above the background noise. Kuschnig et al. (1997) using Monte Carlo simulations established a relationship between the peak height above background and the FAP (see Fig. 4 in their paper). In the case of the $I_{H\alpha}$ amplitude spectrum the peak at $\nu = 0.015 \text{ d}^{-1}$ is 4.3 times above the noise level. This results in $\text{FAP} \approx 10^{-4}$. Both the Fourier amplitude spectrum and the L-S periodogram analysis both indicate that the $I_{H\alpha}$ variations at the orbital frequency of GL 581d are highly significant.

3. $I_{H\alpha}$ versus RV variations

Figure 4 shows the residual $I_{H\alpha}$ variations phased to the orbital period of GL 581d determined by Hatzes (2014). For clarity the data has been phase-binned on intervals $\Delta\phi \approx 0.1$. The error bars represent the standard deviation divided by the square root of the number of points in each bin. The solid curve represents a sine fit to the data. Also shown as a dashed line is the RV orbital solution for GL 581d. The $I_{H\alpha}$ – RV variations for GL 581d are anti-correlated thus supporting the conclusions of Robertson et al. (2014). The variations can be fit by a pure sine function. If the RV variations attributed to GL 581d are actually due to activity, then these should mimic a circular Keplerian orbit. There appears to be a slight phase shift of ≈ 0.1 between the RV and $I_{H\alpha}$ data but this is not deemed significant and may be an artifact of the binning process.

Over the interval JD = 2455200 and 2455645 the sampling of the data (both RV and $I_{H\alpha}$) was excellent and

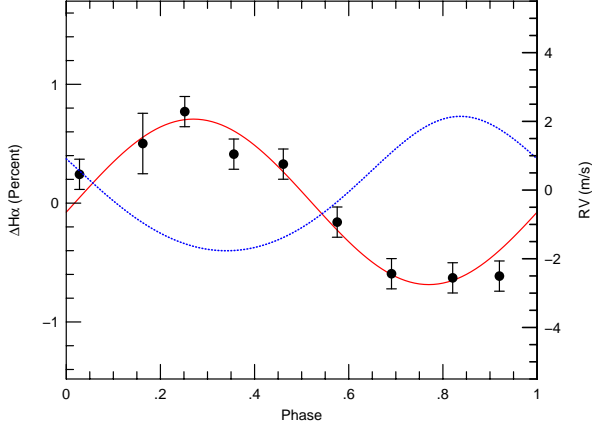


Fig. 4. The phase-binned averages ($\Delta\phi \approx 0.1$) residual $H\alpha$ variations phased to the 125-d period for “GL 581d”. The solid curve represents a sine fit. The dashed curve represents the RV orbital solution for GL 581d.

almost three cycles of variations were covered. The residual $I_{H\alpha}$ data during this time interval are shown in Figure 5 along with the RV variations due to GL 581d (i.e. the contribution of all other planets and activity have been removed). Note that in this case the RV and $I_{H\alpha}$ variations are exactly 180° out-of-phase with each other.

A sine fit to the RV and $I_{H\alpha}$ variations over this interval and allowing the period to be a free parameter resulted in consistent values for the period for the two quantities, namely 71.99 ± 1.51 d and 70.64 ± 1.64 d for the $H\alpha$ and RV, respectively. These periods are slightly longer by about 1.4σ compared to the orbital period of 66.64 ± 0.08 d derived using the full data set. This may be an indication of possible period variations due to differential rotation which may also be consistent with activity-related variations. Unfortunately, we cannot be sure given the short time span of the data.

We also examined the ratio of the amplitudes of the $I_{H\alpha}$ and RV variations ($A_{H\alpha}/A_{RV}$). This ratio should be constant if the two variations stemmed from the same origin. However, we might see temporal variations if the RV amplitude is constant and due to a planet, while the activity signal from $I_{H\alpha}$ varied due to the evolution of the stellar active regions.

Figure 6 shows the ratio $A_{H\alpha}/A_{RV}$ for three different epochs. The $I_{H\alpha}$ amplitude was calculated in parts per thousand (PPT). Although there appears to be a slight decrease in the ratio with time, this is not significant. To within the errors the ratio $A_{H\alpha}/A_{RV}$ is constant with time. However, given that the active regions may evolve at different time scales to those on sun-like stars the constancy of $A_{H\alpha}/A_{RV}$ be less informative due to the restricted time base of the measurements.

4. Discussion

Using a pre-whitening procedure we were able to isolate the variations of $I_{H\alpha}$ in GL 581 at a frequency of $\nu = 0.015$ d $^{-1}$, $P = 66.6$ d which is coincident the RV variations attributed to the planet GL 581d. These variations show sinusoidal variations that are 180° out-of-phase with the

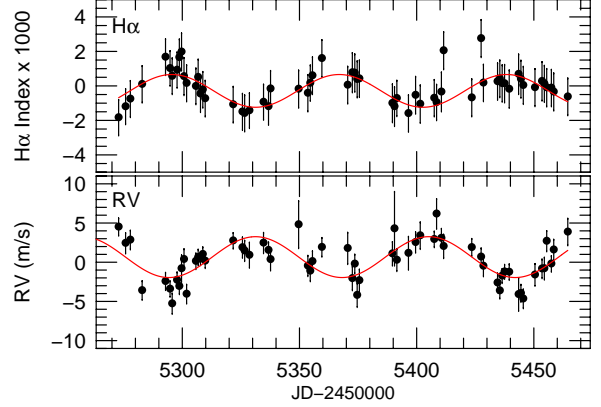


Fig. 5. (Top) The variations in $I_{H\alpha}$ over the time span JD-2450000 = 5272 – 5464. The curve represents a sine fit with the orbital period of GL 581d (Bottom) The RV variations of the purported planet “d” (all other planet signals removed) over the same time span. The curve represents the orbital solution.

Label	Frequency (d $^{-1}$)	Period (d)	Amplitude (PPT)
f_1	0.00719 ± 0.000028	139.16 ± 0.42	0.883 ± 0.079
f_2	0.00797 ± 0.000033	125.39 ± 0.45	0.642 ± 0.071
f_3	0.01517 ± 0.000035	65.91 ± 0.44	0.618 ± 0.063
f_4	0.00037 ± 0.000058	2674 ± 414	0.568 ± 0.100
f_5	0.02282 ± 0.000051	43.83 ± 0.08	0.376 ± 0.006

Table 1. Frequencies found in the $I_{H\alpha}$ data

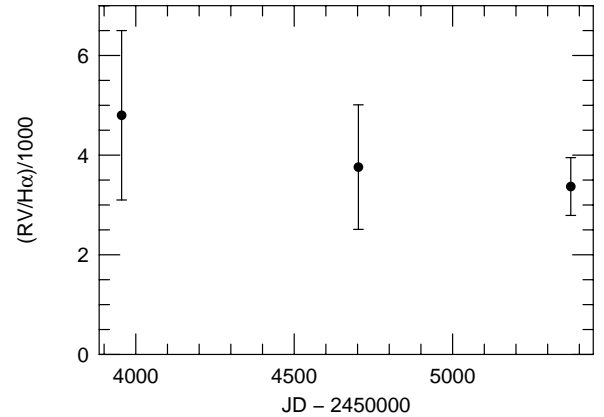


Fig. 6. The ratio of the RV to $I_{H\alpha}$ amplitude ($I_{H\alpha}$ amplitude is in PPT) for three different epochs.

“orbital” RV variations due to GL 581d. This confirms the anti-correlation between the RV variations of GL 581d and $I_{H\alpha}$ found by Robertson et al. (2014). The RV variations attributed to GL 581d are most likely due to stellar activity. GL 581d represents another case where activity-related RV variations can mimic a Keplerian orbit (see Queloz et al. 2001; Bonfils et al. 2007).

The $I_{H\alpha}$ variations appear to be long-lived since they are visible in data spanning almost 7 years. This is certainly

a cautionary tale for RV programs searching for planets in the habitable zone around M-dwarf stars.

Of course, the presence of the period of GL 581d in $I_{H\alpha}$ is no proof that the planet does not exist. There is no physical reason why a planet cannot have the same orbital period as the stellar rotation, or one of its harmonics. There are two arguments against keeping GL 581d as a confirmed exoplanet. First, when one finds significant periodic variations in an activity indicator (photometry, $I_{H\alpha}$, Ca II, bisectors, etc.) with the same period as the RV variations then this is generally accepted as a “non-confirmation” of the planet candidate. There is no reason to change this criterion simply to “save” a habitable planet. It is better to exclude a questionable exoplanet in the census rather than keeping it when it actually is not there.

Second, Robertson et al. (2014) demonstrated that after correcting the RVs due to their correlation with $I_{H\alpha}$ this boosted the significance of the other RV signals that are certainly due to planetary companions. This behavior is also consistent with activity-related RV variations for the 66.7 d signal.

We examined the the ratio of the RV to $I_{H\alpha}$ amplitudes. If the RV amplitude was directly tied to the activity then this ratio should remain roughly constant. On the other hand, if the 66.7 d RV variations were due to a planet then the amplitude of these would remain constant, whereas an activity cycle may produce amplitude variations in the $I_{H\alpha}$ index. This should result in a variation in the RV to $I_{H\alpha}$ amplitude and the planet candidacy of GL 581d could remain. Although our measurements of the RV to $I_{H\alpha}$ amplitude is consistent with no variations, the measurement error is too large to exclude with certainty any temporal variations. Unfortunately, to detect any temporal variations in either the $I_{H\alpha}$ or RV amplitude would require an inordinate amount of additional measurements. At the present time there is no compelling reason to think that GL 581 is more than a 3-planet system.

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